Spectroscopic characterization of 250-µm-selected hyper-luminous star-forming galaxies

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ABSTRACT
We present near-infrared (near-IR) spectroscopic observations from Very Large Telescope Infrared Spectrometer And Array Camera (ISAAC) of 13 250-µm luminous galaxies in the Chandra Deep Field-South, seven of which have confirmed redshifts which average to \(\langle z \rangle = 2.0 \pm 0.4\). Another two sources of the 13 have tentative \(z > 1\) identifications. Eight of the nine redshifts were identified with H\(\alpha\) detection in \(H\) and \(K\) bands, three of which are confirmed redshifts from previous spectroscopic surveys. We use their near-IR spectra to measure H\(\alpha\) linewidths and luminosities, which average to \(415 \pm 20\) km s\(^{-1}\) and \(3 \times 10^{35}\) W (implying SFR\(H\alpha \sim 200\) M\(_\odot\) yr\(^{-1}\)), both similar to the H\(\alpha\) properties of submillimetre galaxies (SMGs). Just like SMGs, 250-µm-luminous galaxies have large H\(\alpha\) to far-infrared (FIR) extinction factors such that the H\(\alpha\) star formation rates (SFRs) underestimate the FIR SFRs by approximately eight to 80 times. FIR photometric points observed from 24 to 870 µm are used to constrain the spectral energy distributions even though uncertainty caused by FIR confusion in the Balloon-borne Large-Aperture Submillimeter Telescope (BLAST) bands is significant. The population has a mean dust temperature of \(T_d = 52 \pm 6\) K, emissivity \(\beta = 1.73 \pm 0.13\) and FIR luminosity \(L_{\text{FIR}} = 3 \times 10^{13}\) L\(_\odot\). Although selection at 250 µm allows for the detection of much hotter dust-dominated hyper-luminous infrared galaxies (HyLIRGs) than SMG selection (at 850 µm), we do not find any \(\gtrsim 60\)-K ‘hot-dust’ HyLIRGs. We have shown that near-IR spectroscopy combined with good photometric redshifts is an efficient way to spectroscopically identify and characterize these rare, extreme systems, hundreds of which are being discovered by the newest generation of IR observatories including the Herschel Space Observatory.

Key words: galaxies: evolution – galaxies: high-redshift – galaxies: starburst – infrared: galaxies.

1 INTRODUCTION
Submillimetre galaxies (SMGs) contribute significantly to the rapid buildup of stellar mass in the Universe at \(z \sim 2\). However, their selection at 850 µm is inherently biased towards colder dust sources (Eales et al. 2000; Blain et al. 2004). Recent work (e.g. Chapman et al. 2004; Casey et al. 2009a) has demonstrated that 850-µm-faint, high-redshift ultraluminous infrared galaxies (ULIRGs) exist and may contribute significantly to the cosmic star formation rate density at its peak. Casey et al. (2009a) describe a population of 70-µm luminous galaxies at \(z \sim 1.5\) whose infrared (IR) luminosities exceed \(\sim 10^{12}\) L\(_\odot\) but are 850-µm faint due to hotter characteristic dust temperatures. Sparse IR data, particularly in the 50–500 µm wavelength range, along with poor volume density constraints have limited the interpretation of these submm-faint ULIRGs. Similar studies of other IR-luminous galaxy populations, selected at 24 µm, 350 µm or 1.2 mm for example, present even more evidence for diverse populations of luminous, dusty starbursts at \(z \gtrsim 1\) which do not necessarily intersect (see Dey et al. 2008; Bussmann et al. 2009; Younger et al. 2009).

The arrival of new IR instruments – including Balloon-borne Large-Aperture Submillimeter Telescope (BLAST; Pascale et al. 2008), Submillimetre Common-User Bolometer Array 2 (SCUBA 2), Large Apex Bolometer Camera (LABOCA) and the Herschel
Space Observatory – has opened up more extensive studies of these distant starbursts. BLAST’s deep mapping of the Extended Chandra Deep Field-South (ECDF-S) at 250, 350 and 500 µm has, for the first time, led to a 2 ULIRG selection near the peak of their spectral energy distribution (SED). Ivison et al. (2010, hereafter I10) and Dunlop et al. (2010, hereafter D10) describe the selection of these 250-µm sources, along with their radio and 24-µm counterparts, in detail and match sources to photometric redshifts derived from the extensive ECDF-S multiwavelength data. We also make use of longer wavelength constraints from the LABOCA 870-µm survey of the ECDF-S (Weiβ et al. 2009). While most low-redshift (z ≲ 0.8) 250-µm sources have spectroscopic identifications, none of the suspected high-redshift sources had spectroscopic redshifts.

This paper presents new Very Large Telescope (VLT) Infrared Spectrometer And Array Camera (ISAAC) spectroscopic observations of 13 BLAST 250-µm sources with zphot > 1. With spectroscopic redshifts, we constrain the far-infrared (FIR) dust SED (implying that they are HyLIRGs with LIR > 10^{12} L⊙), measure dust temperatures, blackbody emissivity, FIR luminosities, Hα luminosities and active galactic nuclei (AGN)/metal lines and constrain the FIR–radio correlation for these high-z galaxies. Throughout, we use a ΛCDM cosmology (Hinshaw et al. 2009) with H_0 = 71 km s\(^{-1}\) Mpc\(^{-1}\) and Ω_m = 0.27.

2 OBSERVATIONS AND RESULTS

Long-slit spectroscopic observations of 250-µm sources were obtained in 2009 November on VLT ISAAC under excellent seeing conditions (0.3–0.7 arcsec in the K band). Spectroscopic candidates were chosen from the I10 and D10 ECDF-S BLAST Deep map samples (rms sensitivity σ_{250} = 11 mJy) with photometric redshifts above z = 1 or undefined photometric redshifts (the latter caused by a lack of high quality photometry). I10 selected sources at >5σ, having folded in the confusion noise (∼21 mJy), resulting in flux densities S_{250} > 59 mJy. D10 selected sources at >3σ without accounting for the confusion noise, so their source list has S_{250} > 33 mJy. All had reliable 24-µm and/or radio counterparts, which were then matched to K-band sources (with offsets of <1 arcsec to the radio/24-µm centroid) in archival Multiwavelength Survey by Yale–Chile (MUSYC) data (Gawiser et al. 2006) for VLT spectroscopic targeting. The resulting candidate object list contained 20 sources.

We refer the reader to D10 and I110 for the analysis of the BLAST 250-µm CDF-S map and source selection, as well as some source properties derived from ancillary data. The two papers present different detection thresholds (which are discussed in more detail as they relate to source density estimations in Section 3.1) and also use different counterpart identification methods to identify sources for photometric redshift fitting. D10 use 24-µm counterpart matching, while I10 use 1.4-GHz radio matching. In general, radio matching is much more reliable assuming that the FIR–radio correlation holds (Helou, Soifer & Rowan-Robinson 1985) as there are far fewer potential counterparts and good reason to suspect that an FIR-bright source is also radio-bright. Identification at 24 µm is a less reliable alternative due to the large IR beamsize and density of sources. We caution the reader that our K-band identifications are nearest neighbours to the radio and 24-µm counterpart astrometry of D10 and I10 and that there is minor potential for misidentified counterparts. The offsets between 250-µm peaks and K-band sources (which are effectively equivalent to the radio/24-µm positions) range from 1 to 16 arcsec, averaging to about 7 arcsec, which is well within the beamsize of 250-µm observations; however, it is not possible to know if the counterparts have been identified correctly without high-resolution FIR observations (e.g. Younger et al. 2010).

We observed 13 of the 20 z_{phot} > 1250-µm sources searching for Hα or [O ii] in J, H and K bands. The band of observations was primarily chosen based on the galaxies’ photometric redshifts (where galaxies with z_{phot} > 2 were observed in the K band and with z_{phot} < 1.8 in the H band, and in the J band for the intermediate region). Galaxies were observed individually under varying seeing conditions which ranged from 0.3 to 0.7 arcsec seeing in the K band. Since all galaxies here are assumed to be unresolved, we varied the slit width according to seeing conditions, minimizing it where possible to reduce skyline contamination. Galaxies were centred on the 2-arcmin long slit and observed in ABBA nodding mode with a 15-arcsec nod. On occasion, two candidates were within a 2-arcmin separation and placed on the same slit, with the maximum possible nod distance, which was sometimes 5–10 arcsec. Data reduction was completed with ESO software combined with IRAF and our own IDL-based routines to obtain 1D and 2D wavelength-calibrated spectra.

Only 13 of the sources were observed due to telescope time constraints. Nine of the 13 have spectroscopic redshifts, seven of which are secure. Three of these seven sources (J033246, J033152 and J033243) were already identified in previous spectroscopic surveys using the VLT/Focal Reducer and Low Dispersion Spectrograph 2 and the Gemini Near-Infrared Spectrograph (Kriek et al. 2008; Vanzella et al. 2008) at z = 1.382, 2.336 and 2.122, respectively. Our measured redshifts confirm these observations. Six of the seven secure redshifts and the two tentative redshifts were measured from Hα detection [at a >4σ Hα signal-to-noise ratio (S/N)]. One of the eight Hα redshifts (J033129) would nominally be tentative, but it was spectroscopically identified in the rest-UV independently; the ninth secure redshift which is not based on Hα (J0333151) has absorption features in the K band at z ~ 1.599 which agree with a rest-UV redshift of 1.605 observed independently (Swinbank, private communication). The lines were identified as Hα using a combination of photometric redshift consistency and a lack of other line features (which would instead identify the line as [O ii] or [O iii], in the case of [N ii] or S ii detection). The four sources which were not identifiable in emission either have very weak emission features or lie at redshifts in the range 1.7 < z < 2.0; Hα at these redshifts falls between H and K bands and is thus not detectable with near-IR spectroscopy.

The nine galaxies have a mean redshift z = 2.0 ± 0.4, and their redshift distribution, with respect to other 250-µm sources, is shown in Fig. 1. We determine that their photometric redshifts [derived in D10 and in Rafferty et al. (in preparation) for I10 sources] are good to dz/(1 + z) ≤ 12 per cent, and we emphasize that this relatively small error implies that near-IR spectroscopic followup for ULIRGs with good photometric redshifts is efficient. Table 1 summarizes the galaxy properties and Fig. 2 shows their ISAAC spectra for those which were spectroscopically identified. The galaxies’ names are derived from their positions in the K band.

The spectra in Fig. 2 are framed around the Hα emission for every source except J033151. Regions where emission lines in the sky’s IR spectrum are significant (with flux densities in excess of ~5 5 s^{-1} nm^{-1} arcsec^{-2} m^{-2}, where 5 represents photons) are masked out in both 1D and 2D spectral renditions. The width of these skylines varied according to the slit width of each observation, which varied from 0.3 to 0.8 arcsec.

1 See the Gemini Observatories IR Background Spectra page (http://www.gemini.edu/?q=node/10787) for example sky spectra.

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We choose not to propagate the uncertainty in the FIR deboosting factor (which is estimated to be as large as 50 per cent at a low 250-µm S/N) into the SED uncertainty because the deboosting factor is not independent between bands. Despite the large uncertainty, the FIR flux densities for individual sources would be deboosted by similar factors. A correlated deboosting factor between bands would imply less uncertainty in derived $T_{\text{dust}}$ or $\beta$ than blindly adopting the deboost uncertainties. We test the correlation of the deboosting factor by using the FIR colours ($S_{250}/S_{350}$ and $S_{250}/S_{500}$) of Herschel Spectral and Photometric Imaging Receiver (SPIRE) and Photodetector Array Camera and Spectrometer (PACS) sources (Amblard et al. 2010). Using Monte Carlo tests, we remove the contribution of a single potential boosting source by subtraction of an arbitrary 250-µm flux and the associated 350- and 500-µm fluxes associated with the colours of a randomly selected galaxy from the Amblard et al. (2010) sample. We find that the FIR luminosity does not vary by more than $\pm 0.1$ dex and that dust temperature varies by about $\pm 9$ K. If two contaminating boosting sources are incorporated with different FIR colour properties, the variance on the fitted $L_{\text{FIR}}$ and $T_{\text{dust}}$ decreases further. We discuss the impact that the deboosting factor uncertainty has on our final conclusions more in Section 3.5.

We measure 870-µm flux densities (at the K-band positions) from the LABOCA map of ECDF-S (Weiβ et al. 2009). Six sources have $S_{870} \gtrsim 4$ mJy and are listed in Weiβ et al. (2009). J033246 is claimed as an SMG in D10, but its 870-µm peak is $\sim 30$ arcsec away from its K-band position. This implies that 7/13 ($\sim 53$ per cent) of our sample are submm-faint and would be excluded from traditional SMG surveys. All galaxies except J033212 are also radio detected in VLA data at $\sim 30$ µJy.

The galaxies’ rest-frame near-IR photometry is consistent with stellar emission, from which we derive stellar masses from the rest-frame H band magnitude (Table 2), using the methods described by Hainline et al. (2009). Measuring the absolute magnitude of a galaxy near its 1.6-µm ‘stellar bump’ provides the most accurate measure of its stellar mass; however, it is reliant on the assumption of a constant mass-to-light ratio (here we assume $M/L = 3.2$), reddening properties and minimal AGN contribution to near-IR flux. The stellar mass estimates are uncertain by $\sim 0.3$ dex. The near-IR photometry is also used to infer AGN content, since a flux excess at 8 µm (significantly above stellar population model fits) is indicative of power-law emission from an AGN. None of our sources has $> 2\sigma$ 8-µm flux excesses.

Two of the 13 observed sources, J033151 and J033152, are X-ray detected above the luminosities which would correspond with their star formation rates ($L_X \gtrsim 10^{44}$ erg s$^{-1}$). Only one individual source shows obvious signs of containing a luminous AGN, from its radio flux excess and detection in the X-rays: J033152. From this AGN estimator (and the analysis in the ensuing section about Hα properties), we infer roughly that 20 $\pm 15$ per cent of 250-µm-bright sources have signs of dominant AGN.

### 2.1 $\text{H}\alpha$ properties

We measure Hα linewidths and [N ii]/Hα ratios in order to infer AGN content from the six BLAST sources for which we have secure Hα observations (i.e. not including J033151). After deconvolving the measured full width at half-maxima (FWHM) with the instrumental resolution measured from skylines in the vicinity of Hα ($\sim 6.5$ Å in the K band and $\sim 4.4$ Å in the H band), we find that our Hα lines have an average rest-frame FWHM of 415 $\pm 20$ km s$^{-1}$ and span the range of 150–800 km s$^{-1}$ (for the six
Table 1. Multiwavelength properties of BLAST 250-µm galaxies.

<table>
<thead>
<tr>
<th>Name</th>
<th>ID</th>
<th>$z_{\text{spec}}$</th>
<th>$z_{\text{phot}}$</th>
<th>$S_{24}$ (µJy)</th>
<th>$S_{250}$ (mJy)</th>
<th>$S_{350}$ (mJy)</th>
<th>$S_{500}$ (mJy)</th>
<th>$S_{870}$ (mJy)</th>
<th>Class</th>
<th>$L_{\text{IR}}$ (10$^{11}$ L$_{\odot}$)</th>
<th>$T_{\text{dust}}$ (K)</th>
<th>$\beta$</th>
<th>$q_{\text{IR}}$</th>
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</thead>
<tbody>
<tr>
<td>Detects</td>
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<tr>
<td>J033129.874–275722.40</td>
<td>J033129</td>
<td>1.482</td>
<td>1.57</td>
<td>270</td>
<td>91.6 ± 11.0</td>
<td>54.4 ± 8.7</td>
<td>46.3 ± 6.2</td>
<td>144 ± 16</td>
<td>SMG</td>
<td>(1.2$^{+0.7}_{-0.4}$)</td>
<td>45.9 ± 3.9</td>
<td>1.2 ± 0.3</td>
<td>2.7 ± 0.5</td>
</tr>
<tr>
<td>J033151.088–274436.91</td>
<td>J033151</td>
<td>1.599</td>
<td>1.91</td>
<td>520</td>
<td>74.0 ± 10.8</td>
<td>63.8 ± 8.5</td>
<td>37.4 ± 5.9</td>
<td>96 ± 13</td>
<td>SMG</td>
<td>(2.2$^{+1.4}_{-0.9}$)</td>
<td>47.1 ± 3.0</td>
<td>1.6 ± 0.3</td>
<td>3.0 ± 0.8</td>
</tr>
<tr>
<td>J033152.090–273926.32</td>
<td>J033152</td>
<td>2.342</td>
<td>2.30</td>
<td>200</td>
<td>78.3 ± 11.0</td>
<td>64.3 ± 8.6</td>
<td>53.1 ± 6.0</td>
<td>965 ± 16</td>
<td>SFRG</td>
<td>(8.1$^{+6.6}_{-4.3}$)</td>
<td>44.5 ± 3.0</td>
<td>2.6 ± 0.3</td>
<td>2.2 ± 0.5</td>
</tr>
<tr>
<td>J033204.849–274647.27</td>
<td>66</td>
<td>2.252</td>
<td>1.94</td>
<td>540</td>
<td>64.3 ± 10.9</td>
<td>62.0 ± 8.4</td>
<td>22.4 ± 6.0</td>
<td>126 ± 12</td>
<td>SMG</td>
<td>(4.0$^{+3.6}_{-2.3}$)</td>
<td>56.7 ± 5.2</td>
<td>1.3 ± 0.4</td>
<td>2.8 ± 0.6</td>
</tr>
<tr>
<td>J033243.209–275514.38</td>
<td>318</td>
<td>2.123</td>
<td>2.09</td>
<td>510</td>
<td>30.1 ± 10.9</td>
<td>32.4 ± 8.5</td>
<td>17.8 ± 6.0</td>
<td>92 ± 10</td>
<td>SMG</td>
<td>(1.4$^{+2.8}_{-0.9}$)</td>
<td>56.9 ± 8.6</td>
<td>0.8 ± 0.7</td>
<td>1.6 ± 0.4</td>
</tr>
<tr>
<td>J033246.329–275327.01</td>
<td>1293</td>
<td>1.382</td>
<td>1.37</td>
<td>200</td>
<td>28.1 ± 10.9</td>
<td>25.3 ± 8.6</td>
<td>14.7 ± 5.9</td>
<td>91 ± 7</td>
<td>SFRG</td>
<td>(0.4$^{+0.1}_{-0.1}$)</td>
<td>53.0 ± 12.9</td>
<td>0.8 ± 0.7</td>
<td>2.5 ± 0.7</td>
</tr>
<tr>
<td>J033249.352–275845.07</td>
<td>J033249</td>
<td>2.326</td>
<td>2.22</td>
<td>320</td>
<td>101.2 ± 10.9</td>
<td>66.4 ± 8.6</td>
<td>22.6 ± 6.0</td>
<td>216 ± 16</td>
<td>SFRG</td>
<td>(8.1$^{+3.9}_{-2.6}$)</td>
<td>56.7 ± 4.5</td>
<td>2.1 ± 0.3</td>
<td>2.8 ± 0.3</td>
</tr>
<tr>
<td>Tentative</td>
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<tr>
<td>J033212.866–274640.89</td>
<td>193</td>
<td>1.93</td>
<td>1.81</td>
<td>40</td>
<td>46.0 ± 10.9</td>
<td>33.2 ± 8.5</td>
<td>8.6 ± 6.0</td>
<td>&lt;40</td>
<td>SFRG</td>
<td>(2.2$^{+4.3}_{-1.4}$)</td>
<td>42.7 ± 9.4</td>
<td>=2.0</td>
<td>&lt;3.20</td>
</tr>
<tr>
<td>J033237.731–275000.41</td>
<td>503</td>
<td>2.64</td>
<td>1.96</td>
<td>210</td>
<td>38.0 ± 10.9</td>
<td>20.0 ± 8.6</td>
<td>16.3 ± 6.0</td>
<td>170 ± 8</td>
<td>SFRG</td>
<td>(3.5$^{+7.5}_{-2.4}$)</td>
<td>46.0 ± 8.4</td>
<td>=2.0</td>
<td>2.4 ± 0.4</td>
</tr>
<tr>
<td>Non-detections</td>
<td></td>
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<tr>
<td>J033221.624–275623.49</td>
<td>158</td>
<td>–</td>
<td>1.85</td>
<td>510</td>
<td>54.9 ± 10.9</td>
<td>28.5 ± 8.4</td>
<td>31.2 ± 6.0</td>
<td>38 ± 8</td>
<td>SMG</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
</tr>
<tr>
<td>J033317.754–274605.96</td>
<td>J033318</td>
<td>2.006</td>
<td>430</td>
<td>79.9 ± 10.8</td>
<td>72.5 ± 8.6</td>
<td>51.4 ± 5.9</td>
<td>100 ± 14</td>
<td>4.3 ± 1.4</td>
<td>SFRG</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
</tr>
<tr>
<td>J033128.792–273916.85</td>
<td>J033128</td>
<td>–</td>
<td>460</td>
<td>105.3 ± 11.1</td>
<td>69.6 ± 8.7</td>
<td>39.8 ± 6.3</td>
<td>35 ± 8</td>
<td>4.5 ± 1.5</td>
<td>SFRG</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
</tr>
<tr>
<td>J033248.011–275416.42</td>
<td>593</td>
<td>&gt; 2.80</td>
<td>&lt;30</td>
<td>18.8 ± 11.0</td>
<td>33.5 ± 8.6</td>
<td>12.2 ± 6.0</td>
<td>44 ± 8</td>
<td>9.3 ± 1.4</td>
<td>SMG</td>
<td>–</td>
<td>–</td>
<td>–</td>
<td>–</td>
</tr>
</tbody>
</table>

*Galaxies are split into three categories: 'detections' (sources which have reliable redshift identifications), 'tentative' (poor quality redshift identifications) or 'non-detections' (no visible emission features). All redshift identifications are based on Hα detection except J033151. The galaxies with 'tentative' identifications are included in all figures and tables of this paper but are excluded from the primary analysis points in Section 3 so as not to affect the interpretation of this paper.

*ID is the identification of the 250-µm source taken from D10 or I10. Those from I10 are of the form J033XXX and correspond to the first half of the BLAST name given in table 1 of I10. Those from D10 appear as two- to four-digit numbers and can be found as the BLAST IDs in table 1 of D10.

*$S_{350}$, $S_{500}$ and $S_{870}$ are measured directly from BLAST/LABOCA ECDF-S maps at their respective wavelengths, using the K-band astrometry for sources in the D10 sample, and 24-µm and 1.4-GHz flux densities are based on nearest neighbour matching (described in I10). The flux densities at 250, 350 and 500 µm have not been corrected for flux boosting, and their uncertainties here only represent instrumental uncertainty; they should be combined in quadrature with the confusion noise (~21 mJy) for an accurate representation of flux uncertainty.

*A galaxy's class is either an SMG or SFRG based on its inclusion as a significant detection in the Weiβ et al. (2009) sample, i.e. if its 870-µm flux density is >24 mJy.

*The ratio of IR luminosity to radio luminosity (as calculated in I10 using $\alpha = 0.75$, see their section 2.2), $q_{\text{IR}}$, is given in the last column.

*Three sources have confirmed redshifts from previous spectroscopic surveys (Kriek et al. 2008; Vanzella et al. 2008).
Figure 2. The VLT ISAAC spectroscopy of nine 250-μm BLAST sources around their Hα emission (0.63–0.69 μm rest wavelength), except in the case of J033151 when no Hα observations were obtained. The first seven sources have secure redshifts, the next two are tentative and the last is a stacked Hα spectrum for the six secure Hα sources. Observations were taken in the H or K band, and both 2D and 1D projections are shown for clarity; scaling is optimized for viewing spectral features which are marked by vertical dashed lines (e.g. Hα, N II). The 2D spectra show an angular scale of 12 arcsec from top to bottom and the 1D spectra are extracted within 0.6 arcsec. Skyline emission features are blocked out (solid grey vertical lines). Higher redshift sources are smoothed more (e.g. as in J033237) since the observed wavelengths (of rest-frame Hα) are higher. The stacked spectrum for the six secure redshift Hα spectra is shown in the bottom right.
Table 2. Hα and stellar properties of BLAST 250-μm galaxies.

<table>
<thead>
<tr>
<th>Name</th>
<th>z</th>
<th>$S_{\text{Hα}}$ (W m$^{-2}$)</th>
<th>$L_{\text{Hα}}$ (10$^{35}$ W)</th>
<th>FWHM$_{\text{Hα}}$ (km s$^{-1}$)</th>
<th>⟨[N II]/Hα⟩$^a$</th>
<th>⟨O/H⟩$^a$</th>
<th>SFR$<em>{\text{Hα}}$ (M$</em>{\odot}$ yr$^{-1}$)</th>
<th>SFR$<em>{\text{FIR}}$ (M$</em>{\odot}$ yr$^{-1}$)</th>
<th>SFR$<em>{\text{radio}}$ (M$</em>{\odot}$ yr$^{-1}$)</th>
<th>$M^*<em>b$ (M$</em>{\odot}$)</th>
</tr>
</thead>
<tbody>
<tr>
<td>J033129</td>
<td>1.482</td>
<td>4.1 × 10$^{-19}$</td>
<td>(5.6$^{+0.9}_{-0.8}$)</td>
<td>199 ± 60</td>
<td>0.13</td>
<td>8.60</td>
<td>2100$^{+1100}_{-700}$</td>
<td>1300$^{+440}_{-130}$</td>
<td>1000$^{+250}_{-150}$</td>
<td>5 × 10$^{10}$</td>
</tr>
<tr>
<td>J033151</td>
<td>1.599</td>
<td>-</td>
<td>-</td>
<td>-</td>
<td>-</td>
<td>3800</td>
<td>1000$^{+320}_{-300}$</td>
<td>3 × 10$^{10}$</td>
<td>1000$^{+320}_{-300}$</td>
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</tr>
<tr>
<td>J033152</td>
<td>2.342</td>
<td>1.3 × 10$^{-18}$</td>
<td>(53$^{+2}_{-1}$)</td>
<td>490 ± 40</td>
<td>0.74</td>
<td>&gt;9.25</td>
<td>419$^{+70}_{-20}$</td>
<td>14000$^{+2200}_{-1400}$</td>
<td>28000$^{+4400}_{-4400}$</td>
<td>3 × 10$^{10}$</td>
</tr>
<tr>
<td>J033204</td>
<td>2.252</td>
<td>5.2 × 10$^{-18}$</td>
<td>(20$^{+1}_{-1}$)</td>
<td>360 ± 40</td>
<td>0.74</td>
<td>&gt;8.64</td>
<td>154$^{+54}_{-21}$</td>
<td>6800$^{+1700}_{-1400}$</td>
<td>3300$^{+1000}_{-800}$</td>
<td>5 × 10$^{10}$</td>
</tr>
<tr>
<td>J033243</td>
<td>2.123</td>
<td>1.3 × 10$^{-18}$</td>
<td>(42$^{+3}_{-2}$)</td>
<td>800 ± 50</td>
<td>0.18</td>
<td>8.71</td>
<td>335$^{+46}_{-10}$</td>
<td>2400$^{+4500}_{-200}$</td>
<td>2000$^{+2200}_{-300}$</td>
<td>4 × 10$^{10}$</td>
</tr>
<tr>
<td>J033246</td>
<td>1.382</td>
<td>4.2 × 10$^{-18}$</td>
<td>(43$^{+10}_{-5}$)</td>
<td>150 ± 70</td>
<td>0.13</td>
<td>8.59</td>
<td>39$^{+4}_{-3}$</td>
<td>680$^{+200}_{-150}$</td>
<td>650$^{+140}_{-140}$</td>
<td>1 × 10$^{10}$</td>
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<tr>
<td>J033249</td>
<td>2.326</td>
<td>5.8 × 10$^{-18}$</td>
<td>(25$^{+3}_{-1}$)</td>
<td>360 ± 50</td>
<td>0.14</td>
<td>8.64</td>
<td>194$^{+20}_{-20}$</td>
<td>14000$^{+7000}_{-3300}$</td>
<td>6100$^{+1900}_{-1300}$</td>
<td>3 × 10$^{10}$</td>
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<td>J033212</td>
<td>1.93</td>
<td>3.5 × 10$^{-18}$</td>
<td>(9.3$^{+1.9}_{-1.6}$)</td>
<td>310 ± 50</td>
<td>0.55</td>
<td>&gt;9.25</td>
<td>73$^{+15}_{-13}$</td>
<td>3800$^{+7200}_{-2000}$</td>
<td>&lt;690</td>
<td>4 × 10$^{10}$</td>
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<td>J033237</td>
<td>2.64</td>
<td>1.0 × 10$^{-18}$</td>
<td>(57$^{+1}_{-1}$)</td>
<td>630 ± 50</td>
<td>0.33</td>
<td>8.96</td>
<td>451$^{+60}_{-60}$</td>
<td>6000$^{+1300}_{-1200}$</td>
<td>6700$^{+1300}_{-1200}$</td>
<td>9 × 10$^{10}$</td>
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</table>

Note. Hα properties from VLT-ISAAC spectra of the BLAST 250-μm sample. The seven from top have secure redshifts (the six which have Hα properties are used in our analysis) while the bottom two have tentative redshifts (we calculate their Hα properties, but exclude them from aggregate property analysis in Section 3 despite being illustrated in figures). FWHM has been deconvolved with the instrumental resolution measured from skylines in the vicinity of Hα (~6.5 Å in the K band and ~4.4 Å in the H band), and SFR is derived from $L_{\text{Hα}}$ using the relation SFR = 7.9 × 10$^{-35}$ $L_{\text{Hα}}$ from Kennicutt (1998).

$^a$The metallicity measurements are computed by $\langle \text{N II}/\text{Hα} \rangle = \log (\text{N II}/\text{Hα})$ and $\langle \text{O/H} \rangle = (\log (\text{O/H}))$ derived from $\langle \text{N II}/\text{Hα} \rangle$ using methods described in Maiolino et al. (2008). The characteristic uncertainty on $\langle \text{N II}/\text{Hα} \rangle$ is ~0.10.

$^b$Stellar masses are measured from Spitzer-IRAC photometry which brackets the rest-1.6-μm stellar bump (see Section 2). The characteristic uncertainty in stellar mass is ~2 × 10$^{10}$ M$_{\odot}$.

For the six galaxies which have secure Hα observations, a stacked spectrum is shown in Fig. 2 which we use to measure the aggregate line emission properties of the sample. The Hα linewidth of the stacked spectrum is 530 ± 280 km s$^{-1}$ (statistically indistinguishable from the individual Hα measurements or the mean SMG linewidth of 390 km s$^{-1}$), and the mean line luminosity corresponds to a star formation rate of 190 M$_{\odot}$ yr$^{-1}$. The linewidth is slightly larger than the mean linewidth for the sample likely due to S/N limitations of the original data. Both linewidth measurements, 415 and 530 km s$^{-1}$, are consistent with the dynamics of active star-forming H II regions, except the high FWHM outlier: J033243 at $z = 2.123$ with FWHM = 800 km s$^{-1}$. J033243 is also the second brightest Hα emitter with a high Hα-implied SFR of 335 M$_{\odot}$ yr$^{-1}$; its SFR$_{\text{FIR}}$/SFR$_{\text{Hα}} = 7$, which is the lowest SFR ratio of the sample indicative of a less Hα obscuration.

We convert the [N II]/Hα ratios to 12 + log (O/H) using the methods described by Maiolino et al. (2008). However, two sources (J033249 and J033212) have O/H limits of <9.25, which correspond to very strong [N II]/Hα (≥0.5). The O/H and [N II]/Hα indicators saturate at metallicities above solar (see Pettini & Pagel 2004), and additional contribution from either AGN or shocked gas can increase the [N II]/Hα ratio further (e.g. van Dokkum 2005). The remaining five secure detections have 12 + log (O/H) values which average to 8.64 ± 0.08. The measured [N II]/Hα ratio for the stacked Hα spectrum implies a metallicity of 12 + log (O/H) = 8.75$^{+0.08}_{-0.10}$ (in agreement with the average for the individual measurements).

2.2 Dust SED fitting and FIR–radio correlation

We fit the MIPS (70 μm, and 160 μm, where available), BLAST (250, 350 and 500 μm), and LABOCA (870 μm) flux densities to two different FIR dust models. For FIR SED fitting, we correct the BLAST flux densities for boosting by confusion noise as mentioned in the beginning of this section. Both FIR SED models assume a modified blackbody emission curve with a single dust temperature:

$$S_{\nu} \propto \frac{\nu^{\alpha + \beta}}{\exp(h\nu/kT_{\text{dust}}) - 1}.$$  

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250-μm galaxies at z>1

Figure 4. The FIR SEDs for high-z BLAST sources, including data from Spitzer-MIPS (24, 70 and 160 μm where available), BLAST (250, 350 and 500 μm) and LABOCA (870 μm). We fit three SEDs to the data: (1) a single dust temperature with fixed $\beta = 2$ modified blackbody as in equation (1) (dash-dotted line), (2) a modified blackbody with $\beta$ treated as a free parameter (solid black line) and (3) an SMG composite SED from Pope et al. (2008) normalized to 24-μm flux density (thin grey line). The galaxies' names, redshifts, best-fitting dust temperatures and FIR luminosities are inset on each SED plot. The two galaxies marked TENTATIVE (from their tentative redshift identifications) are also the two galaxies whose fixed $\beta = 2$ SED fit was significantly better than the $\beta$-free model. Neither tentative galaxies are included in the analysis of Section 3.

where $S_\nu$, the flux density, is a function of rest frequency $\nu$, the emissivity $\beta$, dust temperature $T_d$ and FIR luminosity $L_{\text{FIR}}$ (which governs the normalization of the function). The first model allows $\beta$ to vary (the ‘$\beta$-free’ model) while the second model fixes emissivity to $\beta = 2$. Both models have $T_d$ and $L_{\text{FIR}}$ as free parameters. The advantage of allowing emissivity to vary in the first model allows a reassessment of the emissivity constraints which have been placed on ULIRGs in past studies (e.g. $\beta = 1.5$ or 2.0; Chapman et al. 2005; Casey et al. 2009b,a; Younger et al. 2009). In addition, our measurements of $\beta$ are made independent of any a priori constraint on $T_d$ or $L_{\text{FIR}}$. We choose to make the second model rigid as fits from the first model can be unphysical, as might be the case if the FIR flux densities are particularly faint or affected by source confusion.

Only J033212 and J033237 are poorly fitted to a $\beta$-free model (these are the two galaxies with tentative redshift identifications), since they do not have 70-μm data and have unconstraining upper limits in the FIR. We use only the fixed $\beta = 2$ model for these two. The remaining seven galaxies have reliable $\beta$-free SED fits, and from them we measure $\beta$, $T_{\text{int}}$ and $L_{\text{FIR}}$ (8–1000 μm) (summarized in Table 1). Both fixed $\beta$ and $\beta$-free fits are shown in Fig. 4. We find a mean emissivity of $\beta = 1.73 \pm 0.13$ and a mean dust temperature of $T_d = 52 \pm 6 K$.

The FIR luminosities (8–1000 μm) must be corrected to account for mid-IR (8–25 μm) emission from polycyclic aromatic hydrocarbon (PAH) and power-law sources (e.g. Menéndez-Delmestre et al. 2009) above the single FIR modified blackbody. We tether the Pope et al. (2008) SMG SED to 24-μm flux densities (as seen in Fig. 4) to estimate the luminosity deficit of the single temperature blackbody. This deficit varies substantially object to object due to the large spread in 24-μm flux densities and blackbody properties in the 8–25 μm Wein tail. On average, we find that the contribution of the PAH and AGN emission account for 0.04 ± 0.03 dex of luminosity which we add to the FIR luminosities as a correction factor. Although the mid-IR properties of the sample can vary substantially, this deficit translates to no more than an ∼10 per cent increase in FIR luminosity for these $>10^{12} \text{L}_\odot$ systems. The corrected luminosities are given in Table 1.

We also overplot the composite SMG spectrum, from Pope et al. (2008), normalized to the integrated 24-μm flux density in Fig. 4. While the SMG composite is carefully derived based on mid-IR to FIR data of SMGs to date, it fails to fit the BLAST FIR data on a case-by-case basis. In some cases, it underestimates the FIR luminosities by ±1 dex. This illustrates how a 24-μm-normalized SED fitting procedure, which is common in the literature (e.g. Desai et al. 2009), places poor constraints on the breadth of FIR properties of ULIRG samples, especially in the absence of direct FIR measurements. Recent high-resolution FIR observations (e.g. Younger et al. 2010) have demonstrated that 24-μm counterparts are of-
ten misidentifications and do not correspond to the FIR luminous source.

Although multiple dust temperature blackbodies are found to fit well to local ULIRGs in the literature (e.g. see Clements, Dunne & Eales 2010), strong assumptions must be made regarding the FIR luminosity or normalization, to decompose the sparse FIR data down into multiple blackbody components. Given the uncertainty of the FIR luminosities or flux densities at any given wavelength, we decide to forego multiple dust temperature fitting for well-constrained, single dust temperature blackbody fits. If multiple blackbodies provide a more physical SED fit, then our derived emissivities, from the single blackbody fits, could be underestimated.

We measure $q_{\text{IR}}$, the ratio of integrated IR flux to radio flux, as described in detail by I10 to verify the FIR–radio correlation in our sample. 110 finds a mean $q_{\text{IR}} = 2.41 \pm 0.20$ based on the larger sample of BLAST sources with photometric redshifts. Assuming a radio synchrotron slope of $\alpha = 0.75$, we measure a mean $q_{\text{IR}} = 2.46 \pm 0.18$ which agrees with 110 and earlier findings (e.g. Dale et al. 2007) that there is no evidence for evolution in $q_{\text{IR}}$ with redshift. Since $\alpha$ has a significant impact on the calculation of $q_{\text{IR}}$, we consider the impact of variations in $\alpha$: I10 explicitly measured $\alpha$ for a high-redshift subset of their BLAST sample and found a median value of $\alpha = 0.4$. If we use $\alpha = 0.4$ instead to calculate $q_{\text{IR}}$, we measure $q_{\text{IR}} = 2.27 \pm 0.17$. This is still in agreement with 110’s measurement of the FIR–radio correlation at high redshift within uncertainties.

We note that Kovács et al. (2006) concluded that the local FIR–radio correlation overestimates FIR luminosity by factors of ~0.2–0.4 dex for SMGs, which contrasts with our and I10’s result for BLAST sources. However, the difference is due to different FIR SED fitting procedures; when we re-fit the 21 SMG FIR results for BLAST sources. However, the difference is due to different FIR SED fitting procedures; when we re-fit the 21 SMG FIR results for BLAST sources. However, the difference is due to different FIR SED fitting procedures; when we re-fit the 21 SMG FIR results for BLAST sources. However, the difference is due to different FIR SED fitting procedures; when we re-fit the 21 SMG FIR results for BLAST sources.
in dust-obscured starburst galaxies. However, we note that near-IR spectroscopic observations of SMGs in Takata et al. (2006) measured internal extinction factors of $A_V = 2.9 \pm 0.5$ using Hα/Hβ ratios. When correcting the Hα-inferred SFRs in Table 2 for this dust extinction the FIR-inferred SFRs are recovered, averaging to $\sim 2000 \, M_\odot \, yr^{-1}$. This indicates that dust obscuration is significant in the near-IR and must be corrected for to understand the true nature of the ultraluminous activity in these galaxies.

Placing our [N ii]/Hα metallicity measurements in a larger galaxy evolution context, the metallicities of this sample (measured by converting to (O/H), i.e. $(12 + \log(O/H)) = 8.65$) agree within uncertainties with the observed metallicities of the most massive $z \sim 2$ galaxies, $(12 + \log(O/H)) \sim 8.55 \pm 0.07$, in Erb et al. (2006). The mean [N ii]/Hα ratio for this sample, $0.29 \pm 0.23$, agrees within the uncertainty with the Swinbank et al. (2004) SMGs, $0.41 \pm 0.38$. While evolutionary conclusions should not be drawn from these data alone, the results are consistent with the conjecture that the ULIRG phenomenon occurs at the early stages of a burst in star formation triggered by the merger of two typical gas-rich massive galaxies at $z \sim 2$.  

### 3.3 Temperature fitting and selection

Fig. 6 shows dust temperature ($T_{\text{dust}}$) against FIR luminosity, with BLAST 250-μm sources and SMGs overplotted. Representative $1\sigma$ detection boundaries at 70, 250 and 870 μm are shown to illustrate the populations’ selection biases (2σ, 3σ or 5σ detection limits would have the same shape but be shifted to the right in luminosity; e.g. the 3σ detection limit corresponds to a luminosity shift of $\sim0.4 \, \sigma$). The mean dust temperature of our sample, when fitted with single modified blackbody SEDs, is $52 \pm 6 \, K$, which is comparable to the mean dust temperature of local ULIRGs of similar ($\gtrsim 10^{13} \, L_\odot$) luminosities, $45 \pm 10 \, K$ (Chapman et al. 2003; Rieke et al. 2009), and only $\sim5 \, K$ warmer than SMGs of similar luminosities, $> 10^{13} \, L_\odot$ (and is 15 K warmer than SMGs on average, which are $36 \pm 7 \, K$). Overall, all BLAST sources (except the lower redshift 0J33246) have dust temperatures consistent with the high luminosity end of the SMG distribution.

It is important to note that the SMG FIR luminosities shown here are derived from radio luminosity, via the FIR–radio correlation, and that the associated dust temperature fits are reliant upon that assumption. To first order, we and others have shown that the FIR–radio correlation holds at these redshifts and luminosities (for direct comparison, see Table 2); however, scatter is significant, with differences in measured/derived FIR luminosities of $\pm 0.7 \, \sigma$. As is often done for literature SMGs to date (Chapman et al. 2005), dust temperature is measured by using a single FIR data point (e.g. observed 850 μm) and forcing an SED with fixed radio-inferred $L_{\text{250}}$. If the radio-inferred FIR luminosity is significantly different from

![Figure 6. FIR luminosity against dust temperature for the BLAST 250-μm sample (blue circles, labelled by the last three digits in their right ascension). The dashed ($z=2$) and solid ($z=1$) lines indicate rough $1\sigma$ boundaries at 250 μm (blue, $S_{250} = 11 \, mJy$, BLAST), 70 μm (red, $S_{70} = 1 \, mJy$, MIPS) and 870 μm (black, $S_{870} = 3 \, mJy$, LABOCA), where sources at the given redshift would have $1\sigma$ significance if it has $L_{250}$ and $T_d$ corresponding to the boundary and would be $>1\sigma$ if it lies to the right of the boundary. To translate these curves into 3 σ detection limits, they would be shifted $\sim0.4 \, \sigma$ to the right in luminosity space; the shape of the boundary curves would be maintained. 70 μm-bright ‘hot-dust’ ULIRGs from Casey et al. (2009a) are overlaid as red triangles. The 850-μm-detected SMGs from Chapman et al. (2005) and Kovacs et al. (2006) are overlaid as small black crosses. The area enclosed by the green colour highlights a phase space where 250-μm observations are more sensitive than 870 μm at all redshifts $z > 1$.](http://mnras.oxfordjournals.org/Downloaded from Dalhousie University on May 12, 2016)
the actual FIR luminosity, then the dust temperature will either be grossly over- or underestimated.

We use the BLAST sample and its full SED information (thus directly measuring FIR luminosity) to test the accuracy of FIR fits and derived dust temperature for SMGs and other high-z ULIRGs. Regardless of the accuracy of radio-derived FIR luminosities, we find that $T_{\text{dust}}$ is systematically underestimated by $12 \pm 19$ K when derived from 870-$\mu$m flux densities. Similarly, we also measure dust temperature from the 70-$\mu$m flux densities (as is done in Casey et al. 2009a, for 70-$\mu$m-luminous radio galaxies) and find that they are overestimated systematically by $6 \pm 10$ K.

The severity of these over- and underpredicted dust temperatures is due in part to the sample selection. Because the sample is selected at 250 $\mu$m, it is likely that considering only the 870- or 70-$\mu$m points will produce larger $T_{\text{d}}$ error than the 870- or 70-$\mu$m-selected samples, simply because of the temperature-weighting and biasing of these selection wavelengths. In other words, if a galaxy is 870 $\mu$m bright and 250 $\mu$m faint, then it is far more likely to have a cooler inherent temperature than a galaxy which is bright at both wavelengths. This highlights the difficulty with fitting dust temperatures to single FIR flux measurements and demonstrates that the luminosity–temperature distribution of previously studied ULIRG populations should be revisited when more complete SED information is gathered from Herschel and SCUBA 2 (e.g. Magnelli et al. 2010).

### 3.4 HyLIRG evolution

Fig. 6 highlights where 250-$\mu$m observations are more sensitive to hotter dust temperatures than $\sim$870 $\mu$m at $z > 1$. The sparsity of detections in the highlighted region of Fig. 6 indicates that hot-dust HyLIRGs are genuinely much more rare than their cold-dust analogues at $L_{\text{FIR}} > 10^{13} L_\odot$ at high redshift.

The dearth of hot-dust ULIRGs ($\gtrsim 60$ K) from these data is only significant in the HyLIRG ($>10^{13} L_\odot$) regime for redshifts above $z = 1.5$ (in other words, it is also significant at fainter luminosities at lower redshifts but not fainter luminosities at high redshifts). Due to the 250-$\mu$m BLAST sensitivity, we have only one $10^{12} < L_{\text{FIR}} < 10^{13} L_\odot$ ULIRG in our sample, and its redshift is $\sim 1.3$. This leaves the possibility that $z \sim 2$ hotter dust ULIRGs exist, but lie beneath current 250-$\mu$m imaging depth. Casey et al. (2009a) showed that at slightly lower redshifts, $z \sim 1.5$, star formation-dominated hot-dust ULIRGs ($T_d \sim 52$ K, $L_{\text{FIR}} \gtrsim 2 \times 10^{12} L_\odot$) have been observed at 1/5 the volume density of SMGs, but limitations in Spitzer-MIPS 70-$\mu$m depth prevented significant detections at $z \gtrsim 2$. Also, work by Casey et al. (2009c) argues that hotter dust $> 60$ K HyLIRGs at $z \sim 2$ are less prevalent based on CO observations of submm-faint radio galaxies. After accounting for selection bias, the submm-faint ULIRG sample was $\sim 2 \times$ less luminous in $L_{\text{FIR}}$ and $L_{\text{CO}}$ than CO-observed, cold-dust SMGs.

If we assume a priori that high-z ULIRGs have the same dust temperature distribution as local ULIRGs (which have $T_d = 45 \pm 10$ K above $10^{13}$ K), then there is an $\sim 60$ per cent chance that no $> 60$-K sources are detected within a sample of seven sources (given a Gaussian distribution of dust temperatures for systems of $> 10^{13} L_\odot$). This illustrates how limiting our sample size is when it comes to drawing conclusions for the whole 250-$\mu$m-luminous population. For example, a sample of $\sim 30$ sources with $T_{\text{dust}} < 60$ K must be detected in order for that likelihood to drop to $\lesssim 13$ per cent. More spectroscopic observations and FIR characterizations of similar samples are needed from Herschel and SCUBA 2 for real progress to be made in high-z ULIRG evolutionary studies and to probe the differences with local ULIRG populations.

Despite its significant uncertainty given the small sample size, the lack of hot-dust systems ($> 60$ K) in the BLAST HyLIRG sample is consistent with predictions from smoothed particle hydrodynamics simulations for IR-luminous galaxies (see Narayanan et al. 2009). They suggest that the brightest high-z starbursts ($> 10^{13} L_\odot$) are at their most active phase during the early stages of final merger infall, when gas and dust are diffuse, extended and cold. Warmer dust is suggested to condense either at a later stage merger, when gas and dust have collapsed and heated, or when they have been heated by a growing AGN. While this sample of galaxies exhibits warm dust ($30 < T_{\text{dust}} < 50$ K), we have not detected any $> 10^{13} L_\odot$ hot-dust ($> 60$ K) systems in this sample and find a modest AGN fraction (20 per cent); therefore, our results loosely support the theory that the most luminous HyLIRGs are triggered by major merging events.

### 3.5 Confusion

It is important to once again consider the impact of confusion limitations and deboosting factors on our conclusions. We excluded the uncertainty in the deboosting factor from our results (as discussed in Section 2) because its blind propagation into all of the bands in our SED fits is not justified. As is discussed, the deboosting uncertainty is likely to have far greater effect on $L_{\text{IR}}$ than on $T_d$ or $\beta$ since boosting is correlated between FIR bands. Using a naive model where the main contribution to source confusion is a single additional source within the beam, whose FIR colours have the same distribution as seen in Herschel populations, we estimate that the derived $L_{\text{IR}}$ is uncertain by $\sim 0.1$ dex and $T_{\text{dust}}$ is uncertain by $\sim 9$ K. However, without a proper understanding of the sources which boost the flux of our high-z BLAST sample, it is very difficult to determine how the dust temperatures might change, although it is unlikely that the mean would shift far from the current mean, 52 $\pm$ 6 K. This differential boosting issue should be investigated carefully with future, large Herschel 250-$\mu$m-selected samples.

The deboosting effect on luminosity is easier to quantify than the effect on $T_d$. The maximum uncertainty for the deboosting factor found by Eales et al. (2009) is $\sim 50$ per cent, which would propagate to a factor of $\sim 2$ in luminosity. The mean luminosity of our sample is $\sim 3 \times 10^{13} L_\odot$, a factor of $\sim 5$–10 greater than most high-z ULIRG populations in the literature. The factor of $\sim 2$ difference caused by potential deboosting corrections is not significant in comparison. These 250-$\mu$m-bright galaxies are still ‘HyLIRGs’, thus amongst the most luminous, extreme starbursts measured at high redshift.

### 4 CONCLUSIONS

The redshift identification of these 250-$\mu$m-bright, $z \sim 2$ HyLIRGs has allowed a characterization of their near-IR and FIR properties, leading to the following conclusions.

Near-IR spectroscopy (as probed here by VLT ISAAC) is an efficient way of identifying redshifts (50–70 per cent success rates) for FIR sources which have secure photometric redshifts. The redshift range of our sample is $z = 1.3$–2.6, averaging to $\langle z \rangle = 2.0 \pm 0.4$, making the $z \gtrsim 1$ subset of 250-$\mu$m-bright BLAST galaxies. We also find that H$\alpha$ star formation rates underpredict their FIR SFRs by $\sim 35$ times, and we measure metallicities which are in agreement with other high-z galaxy samples, including those of much lower luminosities but of similar stellar mass.

Having multiple FIR flux densities available for each object, we fit FIR blackbody SEDs to each source and constrain $L_{\text{FIR}}$, $T_d$ and $\beta$.
independent of radio flux density or mid-IR flux densities. We find that the FIR–radio correlation holds, but that SMG composite spectra, when fitted to 24-µm flux densities, do not successfully describe the FIR properties of this sample. We measure FIR luminosities of \(3 \times 10^{13} \, L_\odot\) and dust temperatures averaging \(T_{\text{dust}} = 52 \pm 6 \, K\). However, we warn that both of these conclusions are sensitive to the effects of flux boosting in the FIR, although we estimate that this should not change \(L_{\text{FIR}}\) by more than a factor of \(\sim 2\times\) and \(T_{\text{dust}}\) beyond its quoted error.

Since 250-µm selection is more sensitive to the detection of hotter dust sources than SMG selection (at 850 µm), the lack of >60-K hot-dust galaxies in our sample is potentially an indication that high-z, high-L galaxies are more extended (with diffuse, cool dust) on a whole than local ULIRGs. However, our small sample size limits this conclusion to only \(\lesssim 40\) per cent likelihood. A lack of >60-K hot-dust specimens in the BLAST 250-µm population could be telling to the galaxies’ evolutionary stage; this work highlights the need for more observations of larger samples of similar and fainter sources. FIR mapping at 70–500 µm from Herschel and SCUBA 2 will further select rare and poorly studied high-z ULIRG populations like the galaxies presented here, and near-IR spectroscopic observations will enable further redshift identification of their counterparts, leading to a better characterization of the ULIRG phenomenon at high-z.

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